

次世代宇宙航行と宇宙環境 space propulsion concepts using space environment

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Two types of space propulsion systems using space environment are discussed mainly in terms of orbit control capability.

1. Magnetic Sail in the Solar Wind Flow

Magnetic sail (Magneto-plasma sail) is an interplanetary space propulsion concept which produces the propulsive force due to the interaction between the artificial magnetic field around the spacecraft and the solar wind erupted from the Sun. The thrust can be approximated as continuous outward radial acceleration that is inversely proportional to the square of the radial distance from the sun.

A guidance scheme is proposed for orbital motion under continuous outward radial acceleration that varies in accordance to the solar wind intensity. The maximum attainable radial distance of the outbound trajectory is investigated, and a guidance scheme for achieving this target maximum distance is established under radial acceleration disturbances. The scheme not only provides a control law for continuous radial acceleration but also yields the amount and timing of impulsive maneuvers required to satisfy the guidance requirement at the terminal point.

2. Charged Satellite in the Earth Magnetic Field

The motion of a charged satellite subjected to the Earth's magnetic field is considered. The Lorentz force, which acts on a charged particle when it is moving through a magnetic field, provides a new concept of propellantless electromagnetic propulsion. A dynamic model of a charged satellite, including the effect of the Lorentz force in the vicinity of a circular or an elliptic orbit, is derived and its application to formation flying is considered. Analytical approximations for the relative motion in Earth orbit are obtained. The sequential quadratic programming method is applied to solve the orbital transfer problem. A strategy to reduce the charge amount using sequential quadratic programming is also developed.

キーワード: 宇宙推進, 宇宙環境, 軌道工学

Keywords: space propulsion, space environment, orbital mechanics

深宇宙探査機への宇宙天気研究 -金星探査機あかつきに対する太陽高エネルギー粒子被爆量の評価 II-

The space weather research for deep space probes -Evaluation of solar energetic particles exposure on Akatsuki II-

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Space weather researches have become more and more important, according to the expansion of the "humanosphere" to the space. On the other hand, current space weather researches are mainly for circumterrestrial space, not for the deep space probes that are located far from the earth. We aim to forecast and evaluate the radiation hazard to such space probes far from the earth by using the data taken by the Solar Terrestrial Relations Observatory (STEREO). STEREO provides the images of the part of the Sun that is invisible from the Earth, but only EUV images and coronagraph images are available. First, we examine the possibility of the evaluation of the radiation hazard by using EUV and coronagraph images. It is known that solar energetic particles (SEPs) flux is well correlated with the speed of coronal mass ejection (CME) measured by a coronagraph. We focused on two successive flare/CME events occurred on June 4th, 2011. It occurred in an active region that located on the invisible side of the Sun, and near the disk center as seen from Akatsuki (PLANET-C), the Venus Climate Orbiter that was orbiting the Sun at around 0.7AU. On June 5th, an abrupt decrease in the electric power of Akatsuki was observed, which may be attributed to the effect of SEPs associated with the flare/CME events. We measured the velocity of the two CMEs using the coronagraphic images from STEREO and found that the second CME was much faster (about 2200 km/s) than the first one (about 1000 km/s). Considering the time difference between the two events, it is likely that the second CME caught up the first one before they arrived at 0.7AU. The estimated arrival time is consistent with the timing of the power decrease of Akatsuki. According to a statistical study of CMEs and SEPs performed by Gopalswamy et al. (2004) SEP flux tends to become large if a preceding CME have been launched within 24 hours ahead of the onset time of the primary CME. Using the empirical relationship between the SEP flux and the CME velocity derived by Gopalswamy et al. (2004), we estimate the SEP flux of 10^2 - 10^4 cm⁻² s⁻¹ sr⁻¹. We are also analyzing other large similar events that may potentially affected Akatsuki. We will briefly report these events. We are also analyzing other large events that may potentially affected Akatsuki such as that occurred on January 23 2012.

キーワード: 太陽フレア, CME, SEP, 宇宙天気

Keywords: Solar flare, CME, SEP, Space weather

STEREO Views the Farside: CME and Solar Wind Research from STEREO Observations At large Separation Angles STEREO Views the Farside: CME and Solar Wind Research from STEREO Observations At large Separation Angles

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In the last two years, the STEREO spacecraft have reached opposition and have moved towards the farside of the sun. They are the first man-made probes to image the solar corona and heliosphere from such widely-separated viewing angles. At the same time, we are treated to continuous coverage of the full 360 degree corona during the maximum of solar activity, thanks to the support from AIA and LASCO observations. These unprecedented viewing conditions have opened up new research avenues, such as the study of the lifetime of active regions, the large-scale coupling in eruptive events, and the surprising extent of solar energetic particles in the inner heliosphere.

In this talk, I review the latest research on these areas and outline the future plans for the STEREO observations of the farside of the Sun. I will conclude with a discussion of the implications for Space Weather studies, including an operational Space Weather mission of interest to the Japanese space physics community—a mission to the L5 Lagrangian point.

キーワード: Space Weather, Coronal Mass Ejections, Solar Wind

Keywords: Space Weather, Coronal Mass Ejections, Solar Wind

浮上磁場-コロナアーケード磁場相互作用による太陽プラズマ放出現象のMHDシミュレーションによる研究 Simulation Study of Solar Plasma Eruption by Interaction between Emerging Flux and Coronal Arcade Field

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Many kinds of eruptive phenomena, such as eruptive flares, solar filament eruptions and coronal mass ejections (CMEs) are seen in the solar atmosphere and sometimes have an crucial influence on the earth's magnetosphere. It is widely believed that these eruptive phenomena are caused by the same MHD mechanism by which magnetic energy stored in the corona is released. On the other hand, the detailed mechanism has not been clarified. Many observational studies reported the events which are triggered by the interaction between the newly emerging flux and the coronal arcade field. It is also suggested in these studies that reconnection has an important role in the triggering mechanism. In order to clarify the triggering mechanism by interaction between the emerging flux and the coronal arcade field and its parameter dependence, we perform 2.5-dimensional MHD simulation. The controlled parameters are the magnetic field strength, the location of the emerging flux relative to the arcade field, and the shear angle of the arcade field. As a result, two types of mechanisms are found and these are separated in a parameter space of the location of the emerging flux. One of them appears when the location of the emerging flux is around the polarity inversion line (PIL) of the arcade field and this mechanism depends on reconnection between the emerging flux and the arcade field, as pointed out by the observations. Another appears when the location of the emerging flux is far from the PIL and depends on reconnection in the arcade field above the PIL. We discuss how the interaction between the emerging flux and the arcade field causes eruptions and which kinds of conditions are required. Our results show the possibility to predict whether an eruption occurs or not by investigating the amount of flux of the emerging flux and the distribution of the ambient magnetic field.

キーワード: 太陽フレア, フィラメント放出, コロナ質量放出

Keywords: solar flare, solar filament eruption, coronal mass ejection

Observations from Hinode and SDO of a Twisting and Writhing Start to a Solar-filament-eruption Cascade

Observations from Hinode and SDO of a Twisting and Writhing Start to a Solar-filament-eruption Cascade

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We analyze data from SDO and Hinode of a solar eruption sequence of 1 June 2011 near 16:00 UT, with emphasis on the early evolution toward eruption. Ultimately, the sequence consisted of three emission bursts and two filament ejections. SDO/AIA 304 Ang images show absorbing-material strands initially in close proximity that over ~20 min form a twisted structure, presumably a flux rope with $\sim 10^{29}$ ergs of free energy that triggers the resulting evolution. A jump in the filament/flux rope's height (average velocity ~20 km/s) and the first burst of emission accompanies the flux-rope formation. After ~20 min more, the flux rope/filament kinks and writhes, followed by a semi-steady state where the flux rope/filament rises at (~5 km/s) for ~10 min. Then the writhed flux rope/filament again becomes MHD unstable and violently erupts, along with rapid (> 50 km/s) ejection of the filament and the second burst of emission. That ejection removed field that had been restraining a second filament, which subsequently erupts as the second filament ejection accompanied by the third (final) burst of emission. Magnetograms from SDO/HMI and Hinode/SOT, and other data, reveal several possible causes for initiating the flux-rope-building reconnection, but we are not able to say which is dominant. Our observations are consistent with tether-cutting reconnection initiating the first burst and the flux-rope formation, with MHD processes initiating the further dynamics. Both filament ejections are consistent with the standard model for solar eruptions. NASA supported this work through its Heliophysics program.

キーワード: Sun: CMEs, Sun: filaments, prominences, Sun: flares, Sun: UV radiation

Keywords: Sun: CMEs, Sun: filaments, prominences, Sun: flares, Sun: UV radiation

The science of geomagnetically induced currents The science of geomagnetically induced currents

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Geomagnetically induced currents (GIC) phenomenon impacting long conductor systems on the ground can be considered as the end link of chain of complex physical processes comprising the Sun-Earth system. In this paper I briefly review the current status of our understanding of the physics of GIC and novel applications enabled by the new understanding. More specifically, I will demonstrate how we can follow the chain of physical processes from the solar corona down to the upper mantle of the Earth and to GIC. Further, I will show how state-of-the-art models enable predictive modeling of the entire chain of complex processes.

The potential for severe societal consequences has been driving recent increasing interest in extreme GIC events. I will show how we have addressed the issue by generating 100-year GIC event scenarios. These scenarios are of substantial power grid industry interest and have been fed directly into further engineering analyses. I will review the results of our of 100-year geomagnetically induced current scenarios work and discuss some of the future directions in the field.

キーワード: space weather, geomagnetically induced currents, modeling, extreme events

Keywords: space weather, geomagnetically induced currents, modeling, extreme events

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グローバル MHD シミュレーションを用いた太陽風動圧の大きな急上昇に対する磁気圏応答

Global MHD simulation of the magnetospheric response to large and sudden enhancement of the solar wind dynamic pressure

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New global MHD simulation of the magnetosphere has been developed and is now available for somewhat extreme solar wind input parameters. We started to investigate interesting phenomena including super-large sudden commencement (super SC), geosynchronous magnetopause crossing (GMC), as well as the saturation of cross polar cap potential etc. Here we show the initial results on the response of the magnetosphere to large and sudden enhancement of the solar wind dynamic pressure, gradually increasing the dynamic pressure up to the extreme level of GMC. We discuss the fundamental mechanisms of transient response.

巨大太陽高エネルギー粒子イベントの大気化学への影響 Influence of a giant solar energetic particle event on atmospheric chemistry

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In a giant Solar Energetic Particle (SEP) event associated with a giant solar flare, high-energy protons bombard the stratosphere with high-energy neutrons and high-energy photons. The SEPs ionize and dissociate nitrogen and oxygen molecules along with their tracks, and subsequent chemical reactions induce the change of chemical compositions. Here we investigate the chemical composition changes in the middle atmosphere caused by a giant SEP event. Our calculation includes detailed ion chemistry for the first time to assess this problem.

The SEPs firstly yield charged products (N^+ , O^+ , N_2^+ , O_2^+ , and e^-) and neutral atoms ($N(^4S)$, $N(^2D)$, $N(^2P)$, $O(^3P)$, and $O(^1D)$) in the atmosphere. These ions are relatively short-lived, and are soon converted into neutral chemical species. It is suggested¹⁾, however, that detailed ion chemistry is crucial to understand the chemical composition changes during/after a SEP event. For this, we have considered 12 high-energy radiation processes (radiolysis) and more than 200 chemical reactions as the first step, including ion-molecule reactions, photo-dissociations, ion-ion and ion-electron recombinations, as well as neutral chemical reactions such as NO_x (odd nitrogen radicals), HO_x (hydrogen oxide radicals), and ClO_x (chlorine oxide radicals) cycles. We numerically solve the simultaneous differential equations for these reactions in a reaction network within the so-called box model: a zero-dimensional, *i.e.*, local model where no transport processes are considered. In the box model, climatological temperature for several geometric altitudes between 25 km (lower stratosphere) and 75 km (upper mesosphere) are employed to find the equilibrium state in the concentration of chemical species. The fluence of SEPs is then estimated based on ion-pair production rate calculation²⁾, and is introduced into the equilibrium state in the box model as an energy input perturbation. In this talk, chemical composition changes in the stratosphere during/after the event will be reported on the timescale of several *days*, for which the transport processes are not important. In particular, we will focus our attention on the enhancement in the concentration of reactive nitrogen (NO_y) and on the depletion in the concentration of ozone (O_3) in the stratosphere. This is because NO_y species become the source of nitrate that plays an important role when they precipitate, and because O_3 depletion in the stratosphere changes the temperature and the winds thus may be related with the climate change near the surface.

The results of our box model will be combined with a three-dimensional chemistry-climate model (CCM) to study global influence of the giant SEP event. Here the distribution of chemical compositions is investigated on the timescale of several *years*, where the transport processes are essential and are properly considered. We will then briefly discuss some results from our CCM simulations (in more detail, refer to a presentation by Akiyoshi *et al.* in the session “stratospheric processes and their role in climate” of this meeting).

References

- 1) Funke *et al.*, Atmos. Chem. Phys. **11**, 9089 (2011).
- 2) Jackman *et al.*, Atmos. Chem. Phys. **8**, 765 (2008).

Keywords: solar energetic particles, atmospheric chemistry, model calculations

樹木年輪中 ^{14}C 濃度にみられる突発イベント Rapid events in the carbon-14 content of tree-ring

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Measurement of cosmogenic nuclides, which are radioisotopes produced by cosmic rays, can provide us important information to search past extraterrestrial high-energy events such as supernova, solar proton event (SPE), and so on. Until now, the contents of ^{14}C in tree-rings and ^{10}Be in ice cores have been used for this purpose. However, no clear evidence has been found by ^{14}C and ^{10}Be .

We show the results of ^{14}C content measurement in Japanese cedar annual rings from AD 600 to 1020 with 1- to 2-year resolution, and report two findings of rapid increases of ^{14}C content from AD 774 to 775 and AD 992 to 993. These are clear increases against its measurement errors. The shapes of the two series are very similar, i.e., a rapid increase within one year followed by a decay due to the carbon cycle. The scale of the AD 993 event is 0.6 times as large as the AD 775 event.

The ^{10}Be flux in the Antarctic ice core also shows peaks corresponding to these two ^{14}C events. The proportions of flux increase ($^{14}\text{C}/^{10}\text{Be}$) of the two events are consistent with each other. Therefore, it is highly possible that these events have the same origin.

Although the cause of this event can be explained by a large solar proton event (SPE) or a short gamma-ray burst, we conclude that solar activity is a plausible cause because the occurrence rate of ^{14}C increase events is inconsistent with a observed rate of an short gamma-ray bursts.

磁場なし太陽風での磁気圏電離圏対流

The magnetosphere-ionosphere convection in the solar wind without magnetic fields

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We investigate the magnetosphere-ionosphere convection when merging between the interplanetary magnetic field (IMF) and the magnetospheric one is not effective. To perform the simulation in this condition, we assign the IMF with extremely small intensity. The simulation reveals that the magnetosphere has long magnetotail unlike the closed magnetosphere proposed by Axford and Hines [1961]; this means traditional viscous interaction may not play an important role. It is also obtained that the ionospheric convection exhibits a 2-cell pattern like the magnetosphere-ionosphere convection system for the southward IMF condition. It is revealed that the magnetosphere-ionosphere convection is associated with the R1 field-aligned current generated in the cusp-mantle region where the convection plasma flow traverses the cusp region with enhanced plasma pressure. The traversing plasma flow in the cusp region is fed by the magnetosheath flow which enters into the cusp region along the magnetic field lines. Thus, the magnetosphere-ionosphere stationary convection is performed without the merging.

キーワード: 極端宇宙天気, 太陽風-磁気圏相互作用, 磁気圏電離圏対流, MHD シミュレーション, 複合系, 粘性相互作用

Keywords: Extreme space weather, solar wind-magnetosphere interaction, magnetosphere-ionosphere convection, MHD simulation, compound system, viscous interaction

野辺山電波ヘリオグラフを用いた彩層・コロナ磁場の導出 Measurements of Chromospheric and Coronal Magnetic Fields by Nobeyama Radioheliograph

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太陽大気で生じる諸現象を理解し宇宙天気研究を促進するうえで彩層・コロナの磁場構造は非常に重要である。磁化プラズマ中では、熱制動放射の光学的厚さに左右の円偏波成分間で、磁場強度に比例した差異が生ずる。この原理を用いると、電波の円偏波率と輝度スペクトルの傾きという2つの観測量から視線方向磁場の導出が可能である。本研究では野辺山電波ヘリオグラフによるマイクロ波の円偏波スペクトル観測から、彩層・コロナの視線方向磁場強度の導出を行った。

国立天文台野辺山太陽電波観測所の電波ヘリオグラフは17GHzと34GHzで太陽全面の撮像観測を行い、特に17GHzでは左右両円偏波成分の観測を行っている。本研究では、太陽面の中心付近の活動領域を選択し、17GHzの円偏波成分を解析した。その結果、活動領域で最大数%の円偏波率が存在した。次に17GHzと34GHzの電波輝度の比から電波輝度の周波数スペクトルの傾きを求めた。電波輝度の較正においてはSelhorst et al 2005等の結果を基に17GHz、34GHzの太陽静穏領域の温度をそれぞれ10000K、9000Kと定義し、観測結果を規格化した。この値は静穏領域で約0.15の傾きに相当し、活動領域では0.2から0.6の傾きが観測された。得られた周波数スペクトルの傾き及び円偏波率を用いて磁場強度を求めた結果、対応する領域の光球面磁場強度に対して40から60%の視線方向平均磁場強度が得られた。

得られた視線方向平均磁場強度は彩層磁場とコロナ磁場両方に起因する円偏波信号を含んでいる。本研究では、観測結果をコロナ磁場と彩層磁場に分離するために、光学的に厚い彩層大気と光学的に薄いコロナ大気という2層大気モデルを仮定し、SDO衛星AIAによる紫外線画像、及び光球面磁場をポテンシャル近似したコロナ磁場モデルとの比較を行った。その結果、電波の円偏波信号は活動領域外縁部でコロナ磁場のループ構造と良い対応を示すことが分かった。一方で、活動領域中心では彩層磁場の寄与が強く見られた。以上より、17GHz帯域では彩層・コロナ両層の磁場を導出可能であることが明らかとなった。今後は多周波での円偏波観測により、両層の寄与をより正確に分離することが重要であると言える。

キーワード: 太陽電波, 太陽磁場, コロナ, 彩層, 野辺山電波ヘリオグラフ, 円偏波率

Keywords: Sun: radio radiation, Sun: magnetic fields, Sun: corona, Sun: chromosphere, Nobeyama Radioheliograph, circular polarization degree

Association of Polar Faculae with the Polar Magnetic Patches as Revealed with Hinode Observations

Association of Polar Faculae with the Polar Magnetic Patches as Revealed with Hinode Observations

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Polar faculae are small bright features in the polar region of the Sun. They are observed with concentrations of magnetic fields. Previous studies have shown that the number of polar faculae observed at latitudes greater than 60 degree has 11-year periodicity like the sunspot cycle and that their occurrence maximum is anti-correlated with the solar cycle. The aim of this study is to understand the properties of polar faculae, which are believed to be associated with the polar magnetic patches. We analysed data of the north polar region taken by the Hinode/SOT spectropolarimeter (SP) in September 2007. Accurate measurements of vector magnetic fields at high spatial resolution by Hinode/SP for the first time allow us to compare polar faculae with polar magnetic fields in detail. The continuum intensity map is corrected for limb darkening. There are many patchy magnetic field structures in the polar region and thresholds on both size and intensity for the patches are applied to automatically choose polar faculae pixels. The definition of magnetic patch is same as in Shiota et al. (2012 ApJ). We find that: (1) magnetic patches are not uniformly bright but contain smaller faculae inside. (2) polar faculae tend to have stronger and more vertical magnetic fields than their surrounding within the magnetic patches. (3) almost all magnetic patches with total magnetic flux greater than 10^{18} Mx are associated with polar faculae. From Shiota et al. (2012), such magnetic patches vary with the solar cycle. (4) polar faculae flux accounts only for less than 12% of the total vertical magnetic flux of the polar region observed in this study. But as polar faculae are associated with large magnetic flux concentrations they can be considered as good proxy of the cycle varying component of the polar magnetic field.

Furthermore, we confirm two previous studies: (5) The contrast of polar faculae decreases towards the limb (Okunev and Kneer, 2004). (6) Polar faculae with polarity opposite to that of the polar magnetic field exist, but they are very small in number which is consistent with the result by Blanco Rodriguez et al. 2007.

キーワード: Sun, Photosphere, Polar Magnetic Field, Polar Faculae

Keywords: Sun, Photosphere, Polar Magnetic Field, Polar Faculae

飛騨天文台フレア監視望遠鏡を用いたプラージュインデックスの導出 Derivation of the Solar Plage Index using the Flare Monitoring Telescope at Hida Observatory

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It is well known that the solar irradiance modulates at 11-year solar cycle. There have been many arguments on the relation between the solar cycle and the earth's climate, but at least it is well established that UV/EUV directly affect the upper atmosphere of the earth. The solar UV radiation is mainly emitted by the chromospheric height. In the solar chromosphere, magnetic features such as dark filaments and bright plages are observed. Recently Bertello, Ulrich, and Boyden (2010) developed the Ca II K plage index based on the area occupied by plages and active network on the solar disk, and found good consistency with the UV irradiance.

We aim to derive a proper index of solar UV radiation using the chromospheric H-alpha images observed by the Flare Monitoring Telescope. The Flare Monitoring Telescope has operated at Hida Observatory since 1996 and at Ica University in Peru since 2010. It obtains the full disk images of the Sun at H-alpha center and wings. In this work we analyze the H-alpha center images in order to estimate the area of dark filaments and plages. We report our method to derive the filament and plage indices based on the histogram of the H-alpha intensities. We also make indices separated into the low and high latitudes, to see their differences of their long-term variation.

The preliminary results show that the plage index in the solar minimum 23/24 is lower compared to the one in the solar minimum 22/23. This is consistent with the recent satellite observation of UV radiation (e.g., Stanley et al. 2010). However, if we calculate the plage index only in the high latitude, the indices of the two minima are about the same. Relations between average magnetic field strength, EUV images, and coronal holes will be discussed. Our results tell that the plage index may work as a powerful tool for the estimation of UV radiation back decades in time before the start of the satellite observations, and additionally gives some insights to their origins.

キーワード: 太陽彩層, H α , 紫外線, 太陽輻射量

Keywords: solar chromosphere, H-alpha, ultraviolet, solar irradiance

Long-Term Relativistic Radiation Belt Electron Responses to GEM Magnetic Storms Long-Term Relativistic Radiation Belt Electron Responses to GEM Magnetic Storms

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We present a long-term radiation belt simulation for a 200-day period starting on 25 January 1991, which includes both six geomagnetic storms identified by the Geospace Environment Modeling (GEM) focus group and non-stormy periods of the Combined Release and Radiation Effects Satellite (CRRES) mission, using 3-D time-dependent Versatile Electron Radiation Belt (VERB) code, and compare the simulation results with a multisatellite phase space density (PSD) reanalysis obtained using Kalman filtering of observations from CRRES, GEO, GPS, and Akebono satellites, as well as with the CRRES MEA 1 MeV electron observation. The processes accounted for in the model are radial diffusion-driven by Ultra-Low Frequency (ULF) electromagnetic fluctuations and local (pitch-angle and energy) scattering by plasmaspheric hiss and chorus waves, respectively, inside and outside the plasmasphere. The observations show that a significant decrease in the relativistic electrons in the outer radiation belt is observed in association with the solar wind dynamic pressure enhancement during the main phase of each storm, while during the recovery phase, different types of relativistic electron flux profiles are identified: increased, decreased, and unchanged relative to the pre-storm flux level. First, for an increase of relativistic electrons relative to the pre-storm flux level, the comparison of simulation with reanalysis shows that inward radial diffusion and local acceleration coupled with each other result in a net acceleration. Second, for a decrease or lack of change in relativistic electrons, competing effects of pitch-angle scattering, outward diffusion, and acceleration produce the net decrease in electron PSD and fluxes. The results show that the overall time evolution of the radiation belt is in good agreement with our model simulations, indicating that modeling, including radial diffusion and pitch-angle scattering, is reasonable in predicting the general long-term structure of the outer radiation belt. In addition, with the assistance of local acceleration by chorus waves, the overall flux level in the outer radiation belt becomes comparable to the observation.

キーワード: Radiation Belt, Magnetic Storm, Diffusion Simulation

Keywords: Radiation Belt, Magnetic Storm, Diffusion Simulation

Long-Term Variation of Solar UV/EUV Radiation Examined by Full-Disk Solar Images Long-Term Variation of Solar UV/EUV Radiation Examined by Full-Disk Solar Images

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We report the estimation of long-term variations of solar UV/EUV radiations, which affect on the upper thermosphere, by using full-disk solar images. The SOHO/EIT has shown us full-disk features of the sun in EUVs over 15 years. These data enable us to derive the, spatially resolved, long-term variation of area, brightness of coronal holes, active regions, and so on. In this work we examined the EUV 304 Å emission in different latitudes by using full-disk images taken by SOHO/EIT.

By comparing the EUV 304 Å emissions at solar minima, we found that the EUV emission at low (high) latitude is darker (brighter) at the solar minimum 23/24 than those at the minimum 22/23. We also discuss the relation between the abnormal behaviors at the solar minimum 23/24 and the magnetic field structure by using magnetograms obtained by SOHO/MDI and

by Wilcox Solar Observatory. On the other hand, ground-based chromospheric observations also give us another indicator of solar UV emission, since solar UV radiation mainly comes from the chromosphere. From these data, we try to derive the main features on the solar surface that affect on the upper thermosphere and to estimate the long-term UV/EUV variations.

キーワード: 太陽活動, 長期変動, 紫外線放射

Keywords: Solar Cycle, Long Term Variation, UV/EUV Radiation

高精度全大気圏-電離圏結合モデルの開発 Development of a high-resolution whole atmosphere-ionosphere coupled model

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地球周辺の領域であるジオスペースは、太陽風、磁気圏、電離圏、大気圏からなるが、磁気圏-電離圏-大気圏の領域はお互いに強く結合しており、領域間の相互作用は非線形で極めて複雑である。さらに、太陽光や太陽風の変動、下層大気から伝搬してくる大気波動、大気の組成変化などによって、その振る舞いは一層複雑なものとなっている。このような複雑な系を定量的に理解するには、すべての領域を矛盾無く結合した数値モデルが必要不可欠である。近年、多くの研究機関でさまざまなジオスペースシミュレーションモデルが開発されているが、我々のグループでは、全大気圏-電離圏結合モデル (GAIA: Ground-to-topside model of Atmosphere and Ionosphere for Aeronomy) を開発した。このモデルは、宇宙空間や地上からのさまざまな影響に伴う大気圏・電離圏現象の再現が可能である。我々は大気圏-電離圏相互作用に伴うさまざまな超高層現象の再現と物理過程の解明を進めている。今後、GAIA をさらに高精度化し、次世代の宇宙天気研究／予報ツールとして発展させる予定である。本発表では、これまでの GAIA の成果と、高精度化の計画について報告する。

キーワード: モデル, 大気圏, 電離圏, 結合, 宇宙天気, ジオスペース

Keywords: model, atmosphere, ionosphere, coupling, space weather, geospace

太陽表面下における浮上磁場の検出

Detection of the emerging magnetic flux beneath the visible surface of the Sun

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Solar active regions including sunspots are thought to be the consequence of the emerging magnetic flux from the deeper convection zone. They may cause catastrophic outbursts, i.e., flares and CMEs, into the interplanetary space. Therefore, it is important to study the signature of the emerging flux in the convection zone. However, looking inside the interior by direct optical observations is difficult. In this study, we report the detection of the emerging flux in the uppermost convection zone by helioseismic technique. We use SOHO/MDI Dopplergrams of NOAA AR 10488 to investigate the temporal evolutions of acoustic oscillation powers at six different layers (-15 to -2 Mm) beneath the visible surface of the Sun. As a result, we detect the power reductions up to 2 hours before the flux begins to emerge at the surface. The start times of the power reductions show a rising trend of the order of 1 km s^{-1} , with a gradual deceleration with time. If we assume that the power-reducing agent detected here is actually the magnetic flux, the rising speed of 1 km s^{-1} is well in accordance with previous observations and numerical simulations. The detection of the emerging magnetic flux under the solar surface may allow us to know the mechanism of the magnetic process in the Sun, and may improve space weather science.

キーワード: 太陽, 磁場, 太陽内部, 光球, 宇宙天気

Keywords: Sun, magnetic field, solar interior, photosphere, space weather

Physics-based space weather modeling and forecasting activities at NASA GSFC Space Weather Research Center Physics-based space weather modeling and forecasting activities at NASA GSFC Space Weather Research Center

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NASA GSFC Space Weather Research Center (SWRC) leverages the capabilities of the Community Coordinated Modeling Center (CCMC) that hosts a great number of state-of-the-art physics-based space weather models developed by the international space physics community. In its role to support transition to operations activities, CCMC currently runs many of these models in a real-time mode. The real-time runs enable a wide spectrum of novel space weather products that SWRC uses to provide space weather information to the NASA customer. Examples include real-time execution of heliospheric MHD models that provide ensemble information about background solar wind conditions, predictive simulations of coronal mass ejection (CME) propagation in the interplanetary medium, magnetospheric MHD simulations and real-time modeling of the ionosphere-thermosphere conditions.

In this paper we provide an overview of the current physics-based space weather modeling and forecasting activities at CCMC and SWRC. We will discuss the usage of series of models to cover the entire space weather chain from the solar corona down to the upper mantle of the Earth. Tools such as cone model-based CME analysis procedure that have been developed at CCMC and SWRC to enable efficient usage of the physics-based simulation capacity will be reviewed. Further, CCMC has carried out and supported many verification and validation (V&V) activities that are a core element of model transition to space weather operations. Key elements of these V&V activities are also briefly discussed.

キーワード: space weather, modeling, physics-based simulations, forecasting

Keywords: space weather, modeling, physics-based simulations, forecasting

Dependence of Properties of the Tearing Mode on the Lundquist Number and the Background Environment

Dependence of Properties of the Tearing Mode on the Lundquist Number and the Background Environment

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We perform a set of 2D MHD simulations of the magnetic reconnection process in a long current sheet (CS) developed in the solar eruption, and studied the properties of the CS. The configuration is line-tied to the bottom boundary and open at the other sides. The energy conversion process depends not only on the magnetic Reynolds number, R_m , of the background, but also on the density and the gas pressure inside the CS. The high plasma density and pressure inside the sheet prevents reconnection from taking place quickly. In reality, on the other hand, the CS forms in the disrupting magnetic field and either density or pressure inside may not be high, as suggested by the coronal dimming observed during the early stages of eruptions. This allows fast reconnection to occur in our simulations of solar flares. Further investigations indicate that the time it takes for the first plasmoid to form in the CS increases with the value of R_m , and that the corresponding height decreases.

キーワード: Flares, Magnetic reconnection, Current sheet, Plasma instabilities, Turbulence, Fine structures

Keywords: Flares, Magnetic reconnection, Current sheet, Plasma instabilities, Turbulence, Fine structures

磁気圏尾部の大規模電流層構造の特徴と成因 Characteristics and origin of structured current sheet in the magnetotail

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磁気圏尾部のプラズマシートで形成する電流層の構造とその時間発展は、米国の磁気圏探査衛星群 THEMIS の観測データの解析から明らかになりつつある。電流層内の複数点を同時観測することで、サブストーム成長期には、弱磁場領域、磁場極小領域、分岐した電流層構造が同定されている。これらの構造は、サブストームのメカニズム、磁気リコネクションの物理を調べる上で重要である。本研究では、電流層構造の成因を THEMIS 衛星観測の詳細解析および、3D グローバル電磁流体モデルから調べた。上であげた電流層の構造は、電磁流体モデルでも再現されているが、その成因は必ずしも観測とは同じでない可能性があることがわかった。具体的には、弱磁場領域の形成は、観測でもモデルでも電流層が薄くなる過程と連動しているが、モデルでは観測ほど顕著な減少が起こらない。これは、モデルが断熱圧縮過程で電流層を薄くするのに対して、実際は違うためであると考えられる。発表では、磁気圏尾部における 3D グローバル電磁流体モデルの適用範囲を議論する。

キーワード: サブストーム, グローバルシミュレーション, プラズマシート, 電流層, 分岐した電流層, 弱磁場領域
Keywords: substorm, global MHD simulation, plasma sheet, current sheet, bifurcated current sheet, weak magnetic field

極域地上光学観測による宇宙天気の研究

Doing Space Weather by Using Ground-based Optical Instruments in the Polar Region

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Various ground-based optical observations have been carried out in the polar ionosphere for a long time. Those observations have been done mainly for the purpose of detailed understanding of the plasma physics in the magnetosphere-ionosphere coupling system (physics of aurora etc.). However, such optical data are also helpful for doing space weather studies in the high-latitude part of the Earth. In this talk, we introduce two examples of optical observations in the high-latitude region and discuss those cases in terms of space weather applications in that region.

One example is optical observations of polar cap patches, which are regions high-density plasma in the F region ionosphere streaming from the daytime sunlit region into the dark hemisphere. Ground-based all-sky imager can detect such structures as enhancements of 630.0 nm airglow intensity. Patches are known to be accompanied by smaller scale plasma density irregularities; thus, they would be sources of ionospheric scintillations on the satellite signals in the polar cap region. Within polar cap patches, an enhancement of total electron content (TEC) is also observed, which could be a source of ranging inaccuracy in the satellite-based navigation system such as GPS. In this sense, studies of polar cap patches, in particular understanding of their structuring process (i.e., generation of density irregularities), are fairly important for the space weather applications in the highest latitude part of the Earth.

The other example is ionospheric scintillations on the GPS signal in the auroral latitudes. We often observe such aurora-induced scintillations in the phase of the GPS signal received on the ground at the time of auroral breakup (substorm). An enhancement of TEC of ~10 TECU tends to be observed when the ray path of the GPS signal passes through an intense auroral arc. Such enhancements of phase scintillation and TEC are considered to be manifestations of increase and fluctuation in the electron density at the E region altitude. In this talk, we show how typical GPS receiver responds to the dynamical behavior of breakup aurorae during a relatively large geomagnetic storm, and then demonstrate the impact of the electron density variation in the E region due to the auroral particle precipitations to the satellite-based navigation system.

キーワード: オーロラ, 極冠パッチ, 衛星測位, 電離圏シンチレーション

Keywords: Aurora, Polar cap patches, Satellite Navigation, Ionospheric scintillation

太陽フレア磁気リコネクションの「ひので」分光観測 Hinode Spectroscopic Observation of Magnetic Reconnection in a Solar Flare

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Magnetic reconnection is a key process both in solar flares and substorms in the solar system, so it is important to reveal the fundamental mechanism of onset and energy release via magnetic reconnection in cooperation with solar and space observations. In the Hinode (Solar-B) era, it has become possible to directly measure the plasma flow velocity and electron temperature using the EUV Imaging Spectrometer (EIS). Recently solar activity has increased, and we have succeeded in obtaining spectroscopic observations of a few cusp-shaped flares. The cusp-shape is thought to be the reconnection X-point structure, but whether it is related to slow-shocks or not is still under discussion.

Here we show one of the first spectroscopic observations of a cusp-shaped flare on 2011 February 14. It was simultaneously observed with Hinode, SDO and RHESSI. All these observations enable detailed investigation of the temporal evolution and spatial distribution of the temperature. We found that, first, an emerging magnetic loop triggers magnetic reconnection. This is followed by a hard X-ray burst and a plasmoid ejection. The ejection triggers another neighboring magnetic reconnection, which was observed as a second hard X-ray burst and a cusp-shaped structure. Observed nonthermal electrons are in the energy range of 15-25 keV, and thermal electrons in the cusp-shape are heated up to 10^7 K. This thermal temperature and its distribution were measured by EIS and X-ray telescope (XRT) independently, and their results are almost consistent. During the impulsive phase, we also found bi-directional upward and downward fast flows at the loop-top, i.e. reconnection jet (~ 100 km/s), and a fast evaporation flow at the footpoint of the loops (~ 400 km/s) after some time delay. In this observation, reconnection inflow was too weak to measure compared with the background flows, and the ion temperature was also not successfully measured. These observations would be interesting for future work exploring the details of the slow-shocks attached to the X-point.

Keywords: Magnetic Reconnection, Solar Flare, Hinode (Solar-B)

EUV dimming associated with a coronal jet observed by SDO, Hinode and STEREO EUV dimming associated with a coronal jet observed by SDO, Hinode and STEREO

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We have investigated a coronal jet observed near the limb on 2010 June 27 by the Hinode/X-Ray Telescope (XRT), EUV Imaging Spectrograph (EIS), and Solar Optical Telescope (SOT), and the SDO/Atmospheric Imaging Assembly (AIA), and on the disk by STEREO-A/EUVI. From EUV (AIA and EIS) and soft X-ray (XRT) images we have identified both cool and hot jets. There was a small loop eruption seen in Ca II images of the SOT before the jet eruption. We found that the hot jet preceded its associated cool jet by about 2 minutes. The cool jet showed helical-like structures during the rising period which was supported by the spectroscopic analysis of the jet's emission. The STEREO observation, which enabled us to observe the jet projected against the disk, showed dimming at 195Å along a large loop connected to the jet. We measured a propagation speed of 800 km s⁻¹ for the dimming front. This is comparable to the Alfvén speed in the loop computed from a magnetic field extrapolation of the photospheric field measured 5 days earlier by the SDO/Helioseismic and Magnetic Imager (HMI), and the loop densities obtained from EIS Fe XIV 264.79/274.20 line ratios. We interpret the dimming as indicating the presence of Alfvénic waves initiated by reconnection in the upper chromosphere.

キーワード: Solar activity, Chromosphere, Corona, UV radiation, Spectroscopic, Stereoscopic
Keywords: Solar activity, Chromosphere, Corona, UV radiation, Spectroscopic, Stereoscopic

太陽硬X線観測ロケット FOXSI による高感度撮像分光 High sensitivity hard X-ray imaging and spectroscopy with the Focusing Optics X-ray Solar Imager (FOXSI) sounding rocket

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We made an unprecedented high sensitivity solar hard X-ray (HXR) imaging and spectroscopic observation using focusing optics with the Focusing Optics X-ray Solar Imager (FOXSI) sounding rocket, in collaboration with Space Science Laboratory, University of California, Berkeley, and NASA.

Accelerated electrons in solar flares emit HXRs by bremsstrahlung process as they travel and lose their energy in the solar corona. Therefore, HXR observations of the Sun provide important information about the energy release process in solar flares. Although high sensitivity observations of HXRs from accelerated particles in solar flares are important to investigate solar activity and space weather, Fourier reconstructions are required to obtain images with past HXR instruments such as the RHESSI satellite and the Hard X-ray Telescope onboard the Yohkoh satellite, and the sensitivity is limited. To improve the sensitivity, grazing-incidence HXR focusing optics are a promising new technology for future solar observations. By using focusing optics, arrival directions of incident photons can be determined directly, and image reconstructions are not necessary. FOXSI tested out grazing-incidence HXR focusing optics combined with position-sensitive focal plane detectors for solar observations. The replicated nickel optics are used as the focusing optics, and fine pitch silicon strip detectors with low noise front end ASICs (Application Specific Integrated Circuits) are used as focal plane detectors. In the target energy range of 4-15 keV, the angular resolution of optics is 8 arcseconds with the focal length of 2 m. FOXSI achieves superior sensitivity; two orders of magnitude better than that of RHESSI.

FOXSI was launched on November 2, 2012 and HXR images and spectra from a microflare are successfully obtained. We successfully demonstrated the concept of the high sensitivity instrument, and showed the new vision for the future solar HXR observations. In this presentation, we will report the concept of the FOXSI instrument and observational result. We will also present the plan for the second launch opportunity of the FOXSI rocket, and a spaceborne solar observer featuring similar optics we are to propose.

キーワード: 太陽, 太陽フレア, X線, 粒子加速

Keywords: the Sun, solar flare, X-ray, particle acceleration

磁気嵐に伴う赤道カウンタージェット電流と磁気圏ダイナモ Equatorial counterelectrojets during geomagnetic storms and their possible dynamos in the magnetosphere

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The convection electric field and Region-1 field-aligned currents (R1 FACs) are generated by the solar wind dynamo and are conveyed by the shear Alfvén waves to the polar ionosphere, extending near-instantaneously to low latitude by means of the Earth-ionosphere waveguide. The transmitted electric field drives the DP2 currents at mid-equatorial latitudes with enhancement in amplitude at the dayside equator (EEJ). The convection electric field further extends to the inner magnetosphere and generates the ring current. The partial ring current would work as a dynamo for the overshielding electric field responsible for the equatorial counterelectrojet (CEJ) when the convection electric field reduces its intensity. The convection reduction has been a major cause of the overshielding, but our recent studies [Hashimoto et al., 2011] show that the substorm works as a dynamo for the overshielding electric field even under the steady southward IMF. In this case, the substorm intensifies both the R1 and R2 FACs, but the R2 FACs are strong enough to cause the overshielding. In this paper we show that the CEJ occurs during both the main and recovery phases of the storm and attempt to identify the dynamo for the CEJs. Using solar wind and ground-based data, we suggest that the stormtime CEJ is caused by the northward turning of the IMF, sudden decrease in the solar wind density, and the substorm expansion. It is to be noted that the substorm overshielding plays a crucial role in the generation of the electric field at subauroral latitude responsible for the high speed auroral plasma flow and SAPS.

キーワード: 磁気嵐, サブストーム, 過遮蔽, 赤道カウンタージェット電流, 領域2型沿磁力線電流

Keywords: geomagnetic storm, substorm, overshielding, equatorial counterelectrojet, Region-2 field-aligned current

太陽のグローバルスケールのローカルダイナモ Solar local dynamo in global scale

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We have achieved high resolution calculation of the solar global magneto-convection in spherical geometry with a top boundary at 0.99Rsun. The reduced speed of sound technique (RSST) is adopted in this study. Compared with the anelastic approximation, the RSST has two major advantages: One is the good scaling in parallel computing. Second is the accessibility to the real solar surface. These enable us to use a large grid of 720x1280x3072 and to set the location of the top boundary at 0.99Rsun. Due to large contrasts between the bottom and top regions in the density (>600 times) and the pressure scale height, our calculation includes multi-scale thermal convection, i.e., 100 Mm scale at the base and 10 Mm scale near the top boundary. This type of the small convection pattern is achieved for the first time in the global convection. In our current calculation, the rotation is not included in order to investigate the local dynamo effect in the global scale.

We find that the small scale convection generated near the surface layer unites in relatively deeper layer and becomes strong downflow. Thus at the downflow region, the convection is significantly turbulent and the vorticity has large value. At these regions, the local dynamo is very effective and super-equipartition magnetic field is frequently generated. It is also found that this local dynamo action depends on the location of the surface layer.

ダンジェー対流の誤り Misleading of Dungey's convection

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Magnetosphere-ionosphere (M-I) convection is the magnetohydrodynamic (MHD) motion maintained against the dissipation caused by the ionospheric current system. Primary energy to replenish the ionospheric dissipation is, without any doubt, the solar wind motional energy. In the M-I coupling flow, the magnetospheric motion is transmitted to the ionosphere by the field-aligned current (FAC) to maintain the ionospheric convection potential. The central problem is, therefore, how the solar wind motional energy is converted inside the magnetosphere to the electromagnetic energy of the FAC. Energy conversion process acts between three types of energy, flow motional energy, plasma thermal energy, and the electromagnetic energy. Whereas the general physics ruling these energy conversion is well known, it is not easy to know how, where, and which energy conversion acts in the real configuration of the huge magnetosphere. So that the understanding has not progressed from the cartoons such as the Dungey convection, the Bostrom current, and the current wedge. In recent years, however, the problem is becoming definitely accessible from the development of three dimensional (3D) M-I coupling simulation. Now, we can draw the distributions of energy conversion rate between three kinds of energy, and can trace current lines.

From the simulation results, five electromagnetic energy conversion layers are identified in the magnetosphere, (1) the bow shock, (2) dayside magnetopause-low-latitude cusp, (3) high-latitude cusp-mantle, (4) the plasma sheet, and (5) the inner magnetosphere. The first layer is at the bow shock, where flow motional energy is the driver of the energy conversion. Here we refer the energy associated with the force that does the work for the energy conversion as the driver, and the energy associated with the counter force as the acquire. At the bow shock, most of driving energy (flow motional energy) is converted to thermal energy with some to electromagnetic energy. This configuration is quite reasonable with respect to the nature of the fast shock. The second is at the dayside magnetopause-low-latitude cusp, where magnetic tension released on the dayside is the driver of energy conversion ($J \cdot E > 0$). This driving energy is mainly converted to thermal energy with a little to flow motional energy. The second energy conversion layer on the dayside magnetopause continuously extends to the low-latitude side of the cusp and deposits generated thermal energy to the cusp. Flow motional energy generated in the second layer (it can be called the Dungey flow) is lost and converted to thermal energy through the flow braking, before reaching the cusp. The third energy conversion layer is on the high-latitude side of the cusp to the mantle. Here, thermal energy (driver) is converted to electromagnetic energy (acquire) to generate the region 1 FAC ($J \cdot E < 0$), and equivalently the magnetospheric convection. Strict overlapping is seen between this dynamo layer and the layer in which thermal energy is the driver. These results indicate that the dynamo is driven by the flow crossing pressure gradient and the region 1 FAC is closed by diamagnetic current inside the dynamo.

From the Dungey convection, we first get an image of excitation of convection by the magnetic force. This image must result in the driver layer by electromagnetic energy that overlaps with acquire layer for flow motional energy. If such layer is distributing all over the magnetosphere, it may be just the Dungey convection. In the simulation results, however, such layer distributes only in a narrow region along the dayside magnetopause. The second image from Dungey convection is the dynamo driven by the solar wind motional energy, namely the driver layer by the flow motional energy that coincides with acquire layer for the electromagnetic energy. For this configuration, inertial current will close the FAC. In the simulation results, neither of these two estimation is realized.

キーワード: 対流

Keywords: convection

太陽静穏領域における超粒状斑と磁気ネットワークの起源について On the Origin of the Supergranulation and the Magnetic Network in Solar Quiet Regions

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The supergranulation and the magnetic network are the most conspicuous horizontal structures observed in the lower-atmospheric (UV and IR) emission of the sun.

The origin of these structures remains unsolved for over 50 years. We newly develop a radiative magnetohydrodynamic simulation code for the realistic calculation of the solar surface magneto-convection. The results of two-dimensional convection simulations are presented to show the formation process of supergranulation and the magnetic network.

When the magnetic field strength is moderate, no supergranular peak is found in the kinetic energy spectrum. However, the magnetic energy spectrum has a clear peak at the scale of the magnetic network. The horizontal structure of this magnetic network has a correlation with the horizontal flow at a depth of about 3 Mm. This result is interpreted that the large-scale structure of the magnetic network is formed by the merging of the strong downflows with the smaller scale convection.

When the magnetic field becomes sufficiently strong, the back reaction from the magnetic network to the supergranular convection occurs and the supergranular spectral peak appears in the kinetic energy spectrum. This suggests that the magnetic network is not the result of the supergranulation but the exciter of the supergranular convection in the solar surface.

Based on the results above, we suggest a scenario for the formation process of the magnetic network and supergranulation.

Keywords: supergranulation, magnetic network, photosphere, convection

太陽表面磁極の統計的性質 Statistical properties of magnetic patches on the solar surface

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We investigate statistical properties of magnetic patches on the solar surface, especially velocity and lifetime, by means of auto-tracking algorithm in this study.

The solar surface is covered with the convection and the magnetic field, which are important not only as the energy source of the coronal heating problem but also as the actual example of magneto-convection system on the stellar surface. Despite its importance, the comprehensive description of their properties is difficult because of their complexity and the frequency of their interactions and high frequency of occurrence.

We investigate statistical properties by using our patch tracking method for two sets of line-of-sight magnetograms in quiet regions, one not of which has higher time resolution of nearly 1 minute and the other has long observational period of nearly 140 hours, and our patch tracking method. Both data sets are obtained by the Solar Optical Telescope (SOT) onboard Hinode satellite. We track ~3200 and ~40000 patches as sums of both polarities for each data sets, respectively.

Various relationships are investigated. The distribution of proper velocity of patch structures is investigated. We find that the frequency peak is concentrated around $1.0 \times 10^5 \text{ cm s}^{-1}$ and the median value is $8.0 \times 10^4 \text{ cm s}^{-1}$. We also find that the velocity has a slight negative correlation with flux content of patches and the relation is expressed with a power-law form. The power-law index is derived as -0.23. In addition to the apparent velocity, we investigate the frequency distribution of patch lifetime as well. By comparing the results of the high-cadence data set and the long-duration one, we find the common steep power-law distribution. The power-law index obtained in our analysis is ~-2.5. On the other hand, there is a maximum value in the lifetime around 100 minutes, which is substantially enough shorter than the observational duration.

We will discuss the turbulent diffusion coefficient on the solar surface and the stability of patch structures on the solar surface from these results in the presentation.

キーワード: 太陽, 光球, 磁場, 磁気対流, 自動判別

Keywords: the Sun, photosphere, magnetic field, magneto-convection, automatical detection

PEM05-30

会場:303

時間:5月21日 15:30-15:45

JAXA の宇宙環境計測の現状

Status quo of the JAXA space environment measurement

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In order to monitor space environment and its temporal variations, JAXA Space Environment Group has been developing space radiation detectors and installing them on Low Earth Orbit(LEO) satellites, Geostationary Orbit (GEO) satellites, Geostationary Transfer Orbit (GTO) satellite,

Quasi Zenith Orbit (QZO) satellite and Japanese Experimental Module (JEM) of the International Space Station (ISS).

We are using these space environment data to know the situation of space environment and to provide warning messages to the satellite operators, when the space environment will be harmful. Based on our observation data, we also have constructed a quasi-dynamic radiation belt model for the use in satellite manufacturing.

We report on the status quo of the JAXA space environment measurement.

Solar-Wind Proton Anisotropy Versus Beta Relation Solar-Wind Proton Anisotropy Versus Beta Relation

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Jungjoon Seough^{1*}, Peter H. Yoon¹, Khan-Hyuk Kim¹, Dong-Hun Lee¹

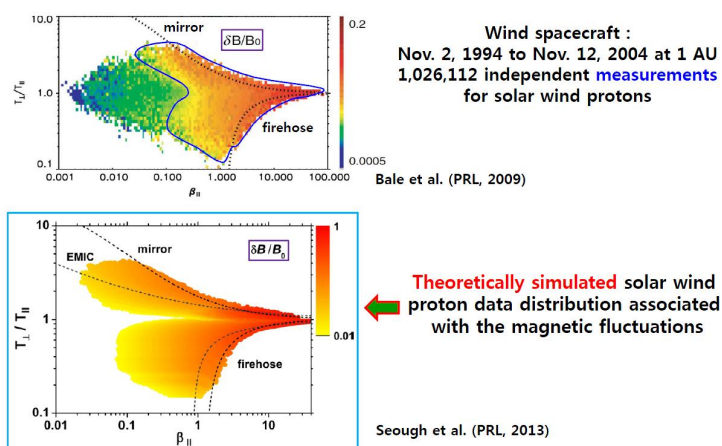
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We address an outstanding problem of the proton temperature anisotropy versus plasma beta inverse correlation in the solar wind. The measured proton temperature anisotropy from the Wind spacecraft at 1 AU is regulated by the oblique (the mirror and oblique fire-hose) instabilities. This observation is inconsistent with the prediction of linear kinetic theory which suggests that the ion-cyclotron and parallel fire-hose instabilities would dominate over the oblique instabilities within a certain range of parallel plasma beta. In the present paper, we put forth a new idea to explain the solar wind observations of the proton anisotropy which do not agree with the current theories. Making use of the fact that the local magnetic field intensity near 1 AU undergoes intermediate-scale temporal variations, we carry out the quasilinear analysis of the temperature anisotropy-driven instabilities with a time-varying local magnetic field, assuming arbitrary initial temperature ratios and parallel betas. It is found that the simulated solar wind proton data distribution in $(\beta_{\parallel}, T_{\perp}/T_{\parallel})$ space is bound by the mirror and oblique fire-hose instabilities, which is superficially similar to the observation.

キーワード: Solar wind proton, Anisotropy-beta relation, Temperature anisotropy-driven instabilities

Keywords: Solar wind proton, Anisotropy-beta relation, Temperature anisotropy-driven instabilities



国際宇宙ステーション日本実験棟曝露部搭載 SEDA-AP での中性子計測 Measurement result of the neutron monitor onboard SEDA-AP on the ISS - Kibo Exposed Facility

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To support future space activities, it is very important to acquire space environmental data related to space radiation degradation of space parts and materials and spacecraft anomalies. Such data are useful for spacecraft design and manned space activity.

SEDA-AP was mounted on Kibo of ISS (International Space Station) to measure the space environment of the 400 kilometres altitude for 3 years.

Neutrons are very harmful radiation because of their strong permeability attributable to its electrical neutrality. The Neutron Monitor measures the energy of neutrons from thermal to 100 MeV in real time using a Bonner Ball Detector and a Scintillation Fiber Detector. The Bonner Ball Detector discriminates neutrons from other charged particles using ³He counters, which have high sensitivity to thermal neutrons. It also measures neutron energy using the relative response, which corresponds to different polyethylene moderator's thickness (6 pcs.). The Scintillation Fiber Detector measures the track of incident particles using a cubic arrangement sensor on which are heaped up 512 scintillation fibers. The sensor discriminates neutrons using differences of these tracks, and measures neutron energy by measuring its track length.

There are three kinds of neutrons measured in space as follows,

1. Albedo Neutron

Caused by galactic cosmic ray and radiation react with atmosphere

2. Local Neutron

Caused by galactic cosmic ray and radiation react with spacecraft

3. Solar Neutron

Caused by accelerated particle in solar flare

Because the shield is difficult, and the influence is large to the human body, the neutron is very important for the astronauts radiation exposure management. Moreover, it is important to measure the albedo neutron because it is thought that the proton that is generated by neutron decay is an origin of the radiation belt. This theory is called as CRAND (Cosmic Ray Albedo Neutron Decay).

An accurate energy spectrum of the solar neutron is measured in space in which the atmosphere is not attenuated, and information on the high energy particle generation mechanism at flare is obtained. This becomes a valuable basic data to do the forecast of flare in the future.

The candidate of some neutron events was found as a result of analyzing the data of the solar flare of M>2 from September, 2009. The detailed analysis on the event on March 7, 2011 was done in these candidates.

This paper reports the development, mission objectives, instrumentation and the result of these analyses.

Keywords: ISS, SEDA-AP

「ひので」による白色光フレア観測 White-Light Flare Observations by Hinode/SOT

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In association with solar flares, we sometimes observe enhancements of visible continuum radiation, which is known as a "white-light flare". These flares are mainly associated with energetic events, such as X-class flares, and they are still only rarely observed since first being discovered more than 150 years ago. Because many observed events show a close correlation between the time profiles and locations of white-light emission, and the hard X-rays and/or radio emission, there is some consensus that the origin of white-light emission is due to accelerated particles, especially non-thermal electrons. During big flares which show white-light emission, huge amounts of electrons are accelerated to high energies ? there might also be huge amounts of protons and ions accelerated to high energies. These large amounts of high energy particles are released into interplanetary space as Solar Energetic Particles (SEPs) and sometimes they reach very near the Earth, and also affect the Earth's environment.

Hinode/SOT has the capability of observing white-light flares in the G-band and continuum (Blue, Green, Red) with a broad-band filter. Using the Hinode Flare Catalogue (Watanabe et al., 2012), we searched for white-light events using G-band and continuum data. We found more than 20 Hinode/SOT white-light events in association with M-class or larger flares between launch (September 2006) and December 2012. We compared the white-light emission data with hard X-ray emission data and/or the strength of the photospheric magnetic fields and looked for any relationship between them.

First, we analyze one of the white-light flares that occurred on December 14, 2006 in detail. We use G-band data from SOT for white light emission and hard X-ray data observed by the RHESSI satellite. We compared the white-light power and the electron power assuming a blackbody for the white light and the thick-target model for the non-thermal electrons, obtaining a good correlation (Watanabe et al., 2010).

However, this is hard to understand in terms of the expected respective emission heights. Theoretically, white-light emission is generated near the photosphere, but non-thermal electrons of energy ~50-100 keV should deposit their energy in the lower chromosphere, more than 500 km above the photosphere.

We investigate this problem with observations of the near-limb X1.7 flare of 27 January 2012, using three continuum bands of the Hinode/SOT. The near-limb location allows us to determine the heights of the emissions. We found the white-light emissions to be located low down, apparently at the photosphere, with the Ca II H emission originating from higher up. We also calculated the temperature distribution from the three white-light continuum bands. The white-light emission temperature is calculated to be about 5500K, and we found that the lower layer has a higher temperature. These findings suggest that high energy particles penetrate to near the photosphere, heating the ambient atmosphere from very low (near photospheric) layers.

In this paper, we present some of the Hinode/SOT white-light events and discuss the flare parameters and origin of the white-light emission.

キーワード: 太陽フレア, 白色光

Keywords: solar flare, white-light

航空機運航中における宇宙線被ばくと宇宙天気情報 Cosmic-ray exposure and Space weather information during aircraft operation

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Effects of exposure to cosmic-ray during aircraft operation are divided into exposure of aircrew and operational impact.

International Commission on Radiological Protection (ICRP) issued a recommendation to include occupational exposure of aircrew with a jet operated exposure from natural radiation source in 1990. Radiation Council consists of the Ministry of Education, Culture, Sports, Science and Technology, the Ministry of Health, Labour and Welfare, the Ministry of Land, Infrastructure, Transport and Tourism established "Guidelines for management of aircrew exposure to cosmic radiation" in 2006. In response to this, airlines keep record of assessed doses on each aircrew using Japanese Internet System for Calculation of Aviation Route Doses (JISCARD-EX) developed by National Institute of Radiological Sciences (NIRS).

Impacts of space weather on aircraft operations can be classified into communications and navigations.

For communication, it includes difficulties on HF radio due to Dellinger Phenomenon while flying out of range of VHF coverages as international flight. And also includes difficulties on SATCOM voice communication and Controller Pilot Data Link Communication (CPDLC) in oceanic region.

Modern navigation by Global Navigation Satellite System (GNSS) is becoming mainstream. GNSS are used all phase of aircraft operation during on the ground, departure, en-route, and approach. Future of operations aim high category precision approach using automatic approach and landing by GNSS even extremely low visibility until stop on runway. Cosmic-ray re-write the data in memory known as soft error on electronic equipment onboard aircrafts.

Use of Space Weather forecast, how to provide the information to aircrew and how to make decisions are urgent consideration.

For these problems International Airways Volcano Watch Operations Group (IAWOPSG) which one of operations group of International Civil Aviation Organization (ICAO) is making draft "Concept of Operations (ConOps) for international space weather information in support of international air navigation". Adoption of ConOps is targeted for ICAO/WMO divisional meeting in 2014.

太陽高エネルギー粒子被ばく線量予測システム WASAVIES の開発:システムの概要 Development of WASAVIES (Warning System of AViation Exposure to SEP): System Overview

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When solar energetic particles (SEP) are incident to the atmosphere, they can induce air showers by generating varieties of secondary particles. Such secondary particles can reach deep into the atmosphere, and enhance the level of radiation doses, which can be a hazard of aircrews. In order to precisely estimate the radiation doses during large SEP events, we are developing a warning system of aviation exposure to SEP, WASAVIES. In the system, the time profile of anisotropic SEP spectrum incident to the Earth are estimated from the latest knowledge of space weather information [1,2], such as basic parameters of flare and coronal mass ejections (CME). The enhancement of the radiation doses at flight altitudes during the event is then calculated from the incident SEP fluxes in combination with a database developed based on air shower simulation performed by the PHITS code [3]. The WASAVIES has been tested and verified by making a comparison between the measured and calculated count rates of several neutron monitors during past GLE (ground level enhancement) events. The final goal of our project is to predict the enhancement of radiation doses due to SEP exposure within 6 hours from the GLE onset.

[1] Kataoka et al. JpGU2013

[2] Kubo et al. JpGU2013

[3] <http://phits.jaea.go.jp/>

キーワード: 太陽高エネルギー粒子, 宇宙線被ばく, 空気シャワーシミュレーション, WASAVIES, PHITS

Keywords: SEP, Cosmic-ray exposure, Airshower simulation, WASAVIES, PHITS

最高エネルギー太陽放射線による航空機被ばく警報システム WASAVIES の研究開発

Development of WASAVIES (WArning System of AVIation Exposure to SEP): Science Modeling

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The prediction of solar energetic particles (SEP) is important to mitigate the space weather hazard toward increasing solar activities, and is also an ultimate problem for physics-based modelers because of the hybrid nature of MHD fluid and particles. We are developing a forecast system called Warning System of AVIation Exposure to Solar energetic particles (WASAVIES). The trigger of WASAVIES is the automated detection of ground level enhancement (GLE) onset by multiple ground-based neutron monitors [Kuwabara et al., Space Weather, 2006]. We then obtain basic parameters of flare and coronal mass ejections (CME) as input parameters for focused transport of SEP [Kubo et al., JpGU2013]. Realistic inner heliosphere is also dynamically reconstructed at the same time [Shiota et al., JpGU2013], and additional control parameters (e.g., mean free path of SEP) are evaluated for solving the focused transport of SEP. Tracing the SEP particles in a Tsytgenenko model, we obtain the time-varying proton rigidity spectra at the top of atmosphere, and the aviation route doses based on the predicted dose-rate are finally evaluated by air shower simulations [Sato et al., JpGU2013]. We show first results and discuss the limitation of the science modeling and possible further development.

Keywords: solar energetic particles, cosmic rays, radiation dose, interplanetary magnetic field, coronal mass ejections

Prediction model for a decay phase of high-energy solar energetic particle events Prediction model for a decay phase of high-energy solar energetic particle events

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Forecasting of solar energetic particle (SEP) events is one of the most important topics in space weather research as SEPs cause severe radiation hazards such as satellite malfunctions, radiation exposure for astronaut, high radiation doses of air crew, and loss of communications by high-frequency radio waves. There are two kinds of SEP forecasting research, forecasting of 1) SEP occurrence itself and 2) time evolution of SEP intensity. In this study, we focus on the time evolution of SEP intensity, especially, for high-energy SEPs having energy of over 100 MeV.

As a typical interplanetary shock can hardly accelerate ions up to 100 MeV, the high energy SEPs are accelerated near the Sun (namely, solar flares and/or coronal shock waves), then transported in interplanetary space followed by being observed at the Earth. We can simulate SEP transport in interplanetary space and reproduce an observation data of high energy SEPs by solving focused transport equation. Therefore, a time evolution of high energy SEP intensity may be predicted by using numerical simulation.

This study is dedicated for a deterministic prediction of decay phase of a high energy SEP event by the numerical simulation and observation data of initial phase of the event.

キーワード: 太陽高エネルギー粒子, 粒子輸送

Keywords: Solar energetic particles, Particle transport

太陽観測データに基づいた内部太陽圏太陽風・CMEの全自動宇宙天気モデルの開発 Development of automatic daily MHD simulation of solar wind and coronal mass ejections in inner heliosphere

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MHD modeling of the solar wind and coronal mass ejections (CMEs) is important to understand the solar-terrestrial environment and to establish space weather forecast because they are the main sources of space weather disturbances. In addition, three-dimensional interplanetary magnetic fields formed by solar wind and CMEs affect the transport of solar energetic particles, and therefore a realistic modeling of the inner heliosphere is required also for the purpose of prediction of solar energetic particles (SEPs) [Kataoka et al., JpGU2013; Sato et al., JpGU2013; Kubo et al., JpGU2013].

We have developed a three-dimensional global MHD simulation of the inner heliosphere. We use daily updated synoptic map of the photospheric magnetic field as a minimal input. As a first step, we calculate coronal magnetic field with potential field source surface model and obtain maps of open magnetic field and expansion factor. Applying empirical models (such as Wang-Sheeley-Arge model), we obtain solar wind synoptic map. Using time series of the solar wind maps as the inner boundary (25 solar radii), we perform the global MHD simulation in 2 AU. The time series of MHD parameters at the Earth position are passed to a radiation belt model [Miyoshi et al. 2004] for forecasting of radiation belt electron flux. These programs are executed everyday on a server in STEL, Nagoya University.

The solar wind as background for propagation of CMEs is prepared in this way. We also report the method to automatically detect flares from observations and to inject associated CMEs, which contains magnetic flux ropes, into the inner boundary of the global MHD simulation. We will introduce modelling results of several CME events associated with high SEP proton flux and discuss the validation of our model and the further developments.

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